

## ASYMMETRIC TRANSIT CURVES AS INDICATION OF ORBITAL OBLIQUITY: CLUES FROM THE LATE-TYPE DWARF COMPANION IN KOI-13

GY. M. SZABÓ<sup>1,2</sup>, R. SZABÓ<sup>1</sup>, J. M. BENKŐ<sup>1</sup>, H. LEHMANN<sup>3</sup>, GY. MEZŐ<sup>1</sup>, A. E. SIMON<sup>1</sup>, ZS. KŐVÁRI<sup>1</sup>,  
G. HODOSÁN<sup>1,4</sup>, ZS. REGÁLY<sup>1</sup>, AND L. L. KISS<sup>1,5</sup>

<sup>1</sup> Konkoly Observatory of the Hungarian Academy of Sciences, P.O. Box 67, H-1525 Budapest, Hungary

<sup>2</sup> Department of Experimental Physics, University of Szeged, Hungary

<sup>3</sup> Thüringer Landessternwarte, 07778 Tautenburg, Germany

<sup>4</sup> Department of Astronomy, Eötvös University, Pázmány Péter sétány 1/A, 1117 Budapest, Hungary

<sup>5</sup> Sydney Institute for Astronomy, School of Physics A28, University of Sydney, NSW 2006, Australia

Received 2011 May 6; accepted 2011 May 24; published 2011 June 24

### ABSTRACT

KOI-13.01, a planet-sized companion in an optical double star, was announced as one of the 1235 *Kepler* planet candidates in 2011 February. The transit curves show significant distortion that was stable over the  $\sim 130$  days time span of the data. Here we investigate the phenomenon via detailed analyses of the two components of the double star and a re-reduction of the *Kepler* data with pixel-level photometry. Our results indicate that KOI-13 is a common proper motion binary, with two rapidly rotating components ( $v \sin i \approx 65\text{--}70 \text{ km s}^{-1}$ ). We identify the host star of KOI-13.01 and conclude that the transit curve asymmetry is consistent with a companion orbiting a rapidly rotating, possibly elongated star on an oblique orbit. The radius of the transiter is  $2.2 R_J$ , implying an irradiated late-type dwarf, probably a hot brown dwarf rather than a planet. KOI-13 is the first example for detecting orbital obliquity for a substellar companion without measuring the Rossiter–McLaughlin effect with spectroscopy.

*Key words:* binaries: visual – brown dwarfs – stars: individual (CCDM J19079+4652AB, KOI-13)

*Online-only material:* color figures

### 1. INTRODUCTION

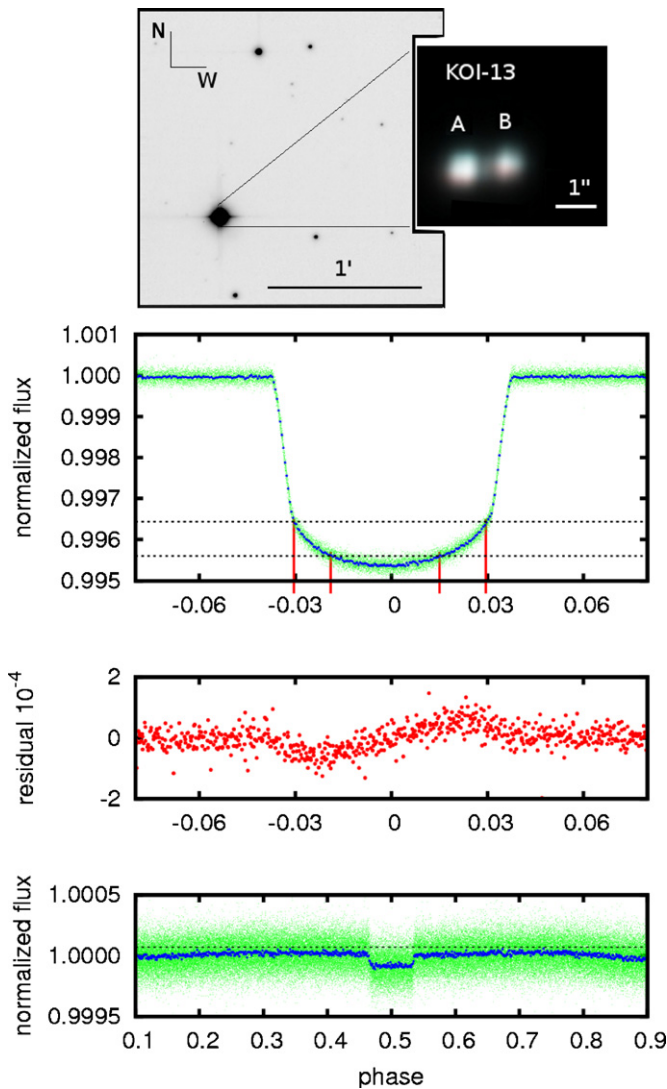
Barnes (2009) suggested that exoplanets orbiting rapidly rotating stars may have unusual light curve shapes. These objects transit across an oblate disk with non-isotropic surface brightness, caused by the gravitational darkening. The von Zeipel (1924) theorem predicts that the flux emitted from the surface of a rapidly rotating star is proportional to the local effective gravity. Therefore, rapid rotation of early-type stars induces an equator-to-pole gradient in the effective temperature. This prediction was recently confirmed by interferometric observations (e.g., Ciardi et al. 2001; Peterson et al. 2006), however, standard models of uniformly rotating stars cannot predict the quantitative findings (Monnier et al. 2007; Che et al. 2011). If a substellar object transits such a star, unusual and distinctive transit light curves are expected. If such asymmetries are measured, one can detect the orbital obliquity of the exoplanet, even without the analysis of the Rossiter–McLaughlin effect in the spectrum (Rossiter 1924; McLaughlin 1924; Gaudi & Winn 2007).

Winn et al. (2010) demonstrated the preference of oblique orbits for hot Jupiters orbiting hot stars ( $T_{\text{eff}} > 6250 \text{ K}$ ). They suggested two alternative hypotheses: (1) the production of hot Jupiters is bimodal: low-mass stars produce a lower number of hot Jupiters which spiral-in the protoplanetary disk, while massive stars form more of massive planets in unstable orbits which then suffer vivid scattering; and (2) there is a single pathway of production of hot Jupiters, but during later evolutionary stages, the interaction between the planet and the convective zone of the star is insignificant for  $T_{\text{eff}} > 6250 \text{ K}$ , which results in the observed bimodality. Currently, there is not enough information on the obliquity distribution of hot brown dwarfs, although one would expect a similar behavior for brown dwarfs and exoplanets in this sense. The discovery of many companions in oblique orbits and the characterization of the host stars are fundamental in choosing between the proposed scenarios.

Borucki et al. (2011) published light curves of 1235 planetary candidates with transit-like signatures detected by the *Kepler* satellite mission between 2 May and 16 September 2009. About 2% of host stars have  $T_{\text{eff}} > 6500 \text{ K}$ , which makes this data set suitable to search for the impact of gravity darkening effect on the measured light curves. We scanned all data, searching for asymmetric transit light curves of early-type stars, and found KOI-13 to be a prime example for such an object. Borucki et al. (2011) remarked that KOI-13 is double, thus a sophisticated analysis was necessary, including follow-up observations. In this Letter, we analyze the distorted light curve of KOI-13 (Section 2), give evidence for the binary nature of KOI-13, and characterize the components (Section 3). Finally, we identify the host star of the transiting compact object and characterize the transiter (Section 4).

### 2. KEPLER LIGHT CURVE OF KOI-13

The visual double star KOI-13 (BD+46 2629, CCDM J19079+4652AB; Figure 1 top panel) was discovered by Aitken (1904), who determined a separation of  $1''.0$  along a position angle of  $281^\circ$ . The companions have apparent magnitudes  $V_A = 9.9 \text{ mag}$ ,  $V_B = 10.2 \text{ mag}$  (CCDM datalog; Dommanget & Nys 1994). The position of the components remained constant since 1904. *Kepler* Short Cadence (SC) Q2 light curve of KOI-13 was taken in the integrated light from both components (Figure 1, upper middle panel), and it shows a significantly asymmetric transit shape. This distortion can be easily perceived by eye. In Figure 1, we plotted the folded SC light curve (epoch = JD 2454953.56498, period = 1.7635892 days; Borucki et al. 2011) with a grid highlighting the most important asymmetries. The lower grid line intersects the light curve at a relative flux of 0.9956, and the orbital phase of the bottom of the transit is marked. Evidently, the floor of the light curve is shifted toward the ingress phase. The flux level of 0.9956 is observed at an orbital phase of  $-0.019$  before the minimum, and 0.015 after



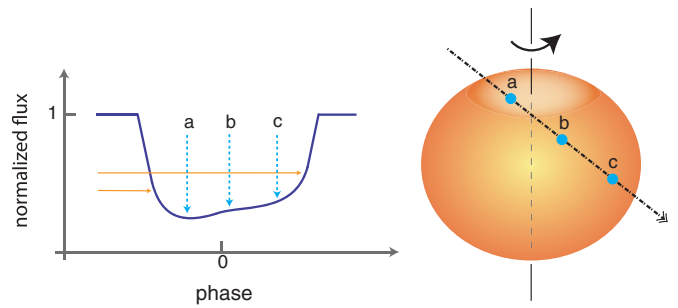
**Figure 1.** Top: the vicinity of KOI-13 (left panel) and its false color ( $V, I$ ) image with lucky imaging (1 m RCC telescope, Konkoly Obs., 2011 April 20). Upper middle: folded *Kepler* light curve of KOI-13 (A and B components together; epoch = JD 2454953.56498, period = 1.7635892 days; Borucki et al. 2011). Lower middle: the residuals of the transit to a symmetric template. Bottom: the out-of-transit phase and the eclipse. Note the grid lines which emphasize the distortions in the light curve.

(A color version of this figure is available in the online journal.)

the minimum, which is a severe light curve distortion. Also, the second contact, which is indicated by a sudden break of the light curve, occurs at 0.9964 flux in the ingress phase (this is pointed by the grid line); while the third contact happens at a slightly increased relative flux, 0.99665; which means 7% difference in the eclipsed light.

Due to these distortions, one can hardly fit any of the widespread symmetric analytical models to the transit curve. The distortions can be checked against a non-parametric template, which is symmetrical in shape, but fits well the observations in other sense. Such a light curve template can be constructed by symmetrization. We constructed it from a spline fit through the phase-folded light curve superimposed with its mirror inversion in phase. The residual curve is plotted in the lower middle panel of Figure 1. The residuals contain about 1/40 of the total light variation.

KOI-13.01 can also be detected in secondary eclipse (Figure 1, bottom panel), with an eclipse depth of  $0.00012 \pm$



**Figure 2.** Observed distortion is qualitatively consistent with stellar temperature gradient due to gravity darkening caused by the rapid rotation of the host star. This schematic figure is based on the Barnes (2009) models.

(A color version of this figure is available in the online journal.)

$0.00001$  and eclipse duration ( $t_1$  to  $t_4$ ) of  $3.0 \pm 0.2$  hr, while its mid-time occurs at  $0.5004 \pm 0.0004$ . There is no hint for an eccentric orbit of KOI-13.01, excluding the eccentricity as the origin of transit shape distortions. The variation of the reflected light is observed in the out-of-transit phases, and it is also asymmetric in shape. The models of Barnes (2009) predict exactly this kind of distortion in the transit light curve if KOI-13.01 ingresses near the hot polar spot of the rapidly rotating star, and the surface is cooler at the end of the transit path (Figure 2).

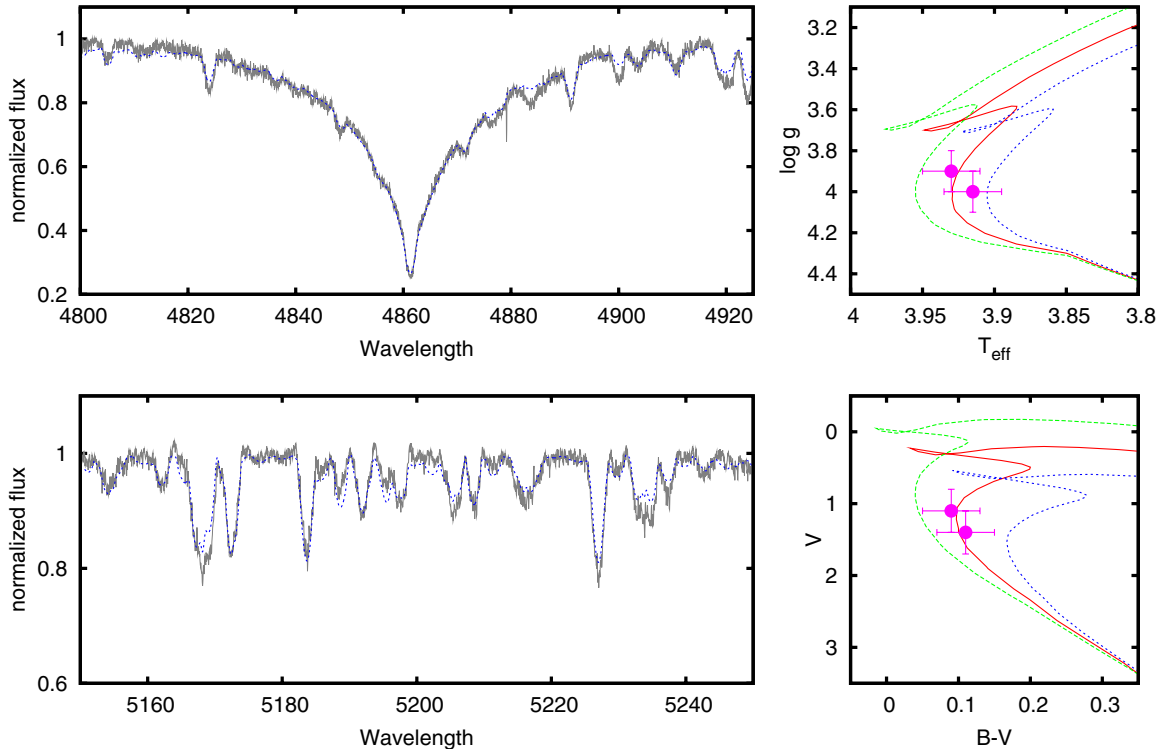
### 3. PROPERTIES OF KOI-13 A AND B

#### 3.1. KOI-13 is a CPM Binary

In the past century, several observations of KOI-13 were taken, including six observations in the *Tycho-2* Catalog. Our latest imaging (Figure 1, top panel) on 2011 April 20 was done at the Pizskéstető Mountain Station of the Konkoly Observatory, with the 1 m RCC telescope equipped with a VersArray 1300b NTE camera (Szabó et al. 2010). The seeing was  $0''.7$  during this observation, enabling the secure separation of the stellar images and an accurate photometry. We determined a separation of  $1''.18 \pm 0''.03$  along P.A.  $281^\circ \pm 1^\circ$ , measured a brightness difference of  $\Delta V = 0.20 \pm 0.04$  mag and almost identical colors,  $\Delta(V - I) = 0.10 \pm 0.05$  mag.

There are four independent positional measurements from the past century, published in Aitken (1904), Dommanget & Nys (1994), Høg et al. (2000), and Mason et al. (2001). It is evident that the position angle of KOI-13 A and B remained constant with a value of 281 within  $2^\circ$ , with no variation in time. Separation values were determined between  $1''.1$  and  $1''.15$ , which is constant or may have been very slightly increasing. Interstellar reddening is significant toward KOI-13: despite both components have an A-type spectrum, *Tycho* ( $B_T - V_T$ ) colors of components A and B are 0.228 and 0.256, respectively.

Two proper motion determinations are available for both components. *Tycho-2* Catalog gives a common proper motion (CPM) for both components,  $\mu_\alpha = 1.5 \pm 1.5$  mas yr $^{-1}$  and  $\mu_\delta = -16.6 \pm 1.4$  mas yr $^{-1}$ . Most recently, Kharchenko & Roeser (2009) determined two independent proper motions,  $\mu_\alpha^A = -0.82 \pm 1.5$ ,  $\mu_\delta^A = -16.30 \pm 1.4$ ,  $\mu_\alpha^B = -1.50 \pm 1.5$ ,  $\mu_\delta^B = -16.60 \pm 1.4$ . This is also consistent with a CPM within the errors. The direction of the proper motion is  $185^\circ \pm 6^\circ$ , which is almost perpendicular to the relative position of the components. Thus, a significant difference in the velocity of proper motion would largely change the position angle. A variation of about  $4^\circ$ – $5^\circ$  in the position angle surely would have been detected within the past years. From the lack of this



**Figure 3.** Determination of stellar parameters of KOI-13. Left panels: details from the observed spectrum (gray) with the best-fit model (dashed, blue online). Right panels: fitting of age, mass,  $B - V$ , and luminosity with isochrones. Lines show the probed Padova isochrones with  $\log t = 8.75, 8.85, 8.95$ , respectively.

(A color version of this figure is available in the online journal.)

observation, we get an upper limit for the *difference* of the proper motions, which is  $S \times \tan 4^\circ / 106 \text{ yr} = 0.7 \text{ mas yr}^{-1}$  (here  $S = 1100 \text{ mas}$ , the separation of the components).

We checked how likely neighboring stars can reproduce the proper motion of KOI-13 and mimic a CPM binary by chance. We extracted data for 151 stars from the *Kepler* field from *Hipparcos* or *Tycho* catalogs, with parallaxes between 1.78 and 2.27 mas. Their mean proper motion is  $\mu_\alpha = 0.97 \text{ mas yr}^{-1}$ ,  $\mu_\delta = -2.8 \text{ mas yr}^{-1}$  with a large standard deviation (10.6 and 11.8, respectively). The ambiguity in the velocity of KOI-13 B is less than  $0.7 \text{ mas yr}^{-1}$ . Only 5% of the stars in its vicinity exhibit a proper motion within this limit around the measured value of  $16.6 \text{ mas yr}^{-1}$ . The direction of P.A.  $45^\circ/225^\circ$  is preferred the most by *Kepler* field stars; proper motion of KOI-13 is inclined to this angle. Only 1%–2% of the stars exhibit a proper motion toward P.A.  $180^\circ\text{--}190^\circ$  and we infer that less than 0.1% of the stars exhibit similar proper motion vector than KOI-13. In conclusion, kinematic information confidently suggest that KOI-13 A and B form a binary system.

### 3.2. The Stellar Components

We fitted stellar model spectra to the A and B components making use of the available photometry and a spectroscopic observation. We took one high-resolution ( $R \approx 32,000$ ) optical spectrum at the Tautenburg Observatory on 2011 April 19, using the 2 m Alfred Jensch telescope and the Coudé echelle spectrograph. The exposure time was 30 minutes, yielding an S/N of about  $125 \text{ pixel}^{-1}$ . Binarity was not seen during the observation because of the seeing, and the slit was centered on the photocenter. The observed spectrum therefore includes the light from both components. In addition, we derived (reddened)

$(B - V)_A = 0.222$  and  $(B - V)_B = 0.227$  from *Tycho* colors. We also made use of the Two Micron All Sky Survey (2MASS) photometry of KOI-13,  $K = 9.425 \text{ mag}$  (the two components together), leading to a reddened  $(V - K)_{A+B} = 0.52$  and  $(J - K)_{A+B} = 0.041$ . Some of these data (the spectrum and  $V - K$  color) were derived from the composite light of both components, only the brightness and  $B - V$  were measured for each component separately. The derivation of stellar parameters was therefore a complex task.

First we fitted the average  $T_{\text{eff}}$ ,  $\log g$ , and  $v \sin i$  of the composite spectrum, assuming solar metallicity. Fixing these parameters and varying metallicity, we determined the metallicity to be  $[\text{Fe}/\text{H}] = 0.2$ , which we assumed for both stars. Then we calculated composite model spectra from atmosphere models with LLmodels program (Shulyak et al. 2004) in its most recent parallel version, and the SynthV program (Tsymbal 1996) to compute the grid of synthetic spectra on a cluster computer at TLS. Based on our lucky image photometry, we assumed 0.55 and 0.45 relative flux of the two individual components, and re-fitted radial velocity and  $v \sin i$  values for each star in the region of metallic lines between 5000 and 5700 Å. A rapid stellar rotation has been revealed for both components, with  $(v \sin i)_A = 70 \pm 10 \text{ km s}^{-1}$  and  $(v \sin i)_B = 65 \pm 10 \text{ km s}^{-1}$ . Because no significant difference was detected in the fitted radial velocities, we took the measured  $-7 \text{ km s}^{-1}$  radial velocity for both components.  $T_{\text{eff}}$  and  $\log g$  were then fitted using the entire 4700–5700 Å wavelength range, including the  $H\beta$  region. A  $\log g$  of 3.9 and 4.0, respectively, resulted for the two components, and we experienced slightly larger  $\log g$  values for the fainter component. The final model spectrum fit is shown by the dashed (blue in the online version) curve in the left panels of Figure 3.

The measured value of  $\log g$  suggests that both components are evolved turnoff stars. Thus, the determination of age, masses,



**Table 1**  
Fitted Stellar Parameters of KOI-13 A and B

Parameter	Comp. A	Comp. B
$\log T_{\text{eff}}$	$3.930 \pm 0.020$	$3.915 \pm 0.020$
$\log g$	$3.9 \pm 0.1$	$4.0 \pm 0.1$
[Fe/H]	0.2	0.2
$v \sin i$	$65 \pm 10$	$70 \pm 10$
$B-V$	$0.09 \pm 0.04$	$0.11 \pm 0.04$
$M/M_{\odot}$	2.05	1.95
$R/R_{\odot}$	2.55	2.38
$M_V$ (mag)	1.06	1.35
$A_V$ (mag)	0.34	0.34
$\log t$	$8.85 \pm 0.1$	$8.85 \pm 0.1$
Dist.	500 pc	500 pc

and reddening is possible from isochrones. For this task, we used Padova isochrones (Bertelli et al. 2008) with  $Z = 0.016$  metallicity. Since the turnoff is a good indicator of age, first we fitted the age of the appropriate isochrone in the  $T_{\text{eff}}-\log g$  plane, which is unaffected by interstellar reddening. The best-fit value was  $\log t = 8.85$  (700 Myr). Then the initial masses of both components,  $E(B - V)$ ,  $E(V - K) = 2.69 \times E(B - V)$  (Larson & Whittet 2005) and  $A_V = 3.1 \times E(B - V)$ , were fitted simultaneously, observing the difference between the model-predicted and the measured  $T_{\text{eff}}$ ,  $\log g$ , and dereddened colors. Finally, we determined two A-type stars with  $L = 30.5 L_{\odot}$  and  $23 L_{\odot}$ , both evolving to subgiant state and rotating rapidly (Figure 3). The fitted parameters and their confidence regions are summarized in Table 1.

### 3.3. The Host Star of the Transiter: KOI-13 A

KOI-13 is a double star, but it is unresolved in the *Kepler* images, because of the coarse image sampling of  $3''.98 \text{ pixel}^{-1}$ . The host star of KOI-13.01 has to be identified before further investigation. With two independent tests, we found that that the brighter star, KOI-13 A, hosts the transiting component.

We did the pixel-level photometry based on the publicly available *Kepler* target pixel files.<sup>6</sup> Few pixels in the central column of KOI-13 image are saturated, but the neighboring 18–20 pixels can be evaluated in each segment in the wings of KOI-13 images. We segmented the vicinity of KOI-13 into four quarters, along the pixel's  $X$ - $Y$  direction. Because the *Kepler* field is rotated, KOI-13 A is in the left segment, and the largest light variation of KOI-13 A is expected to be detected here. The transit depth was measured to be 7.1 mmag in the left image quarter, while the depth was 3.8–5.7 mmag elsewhere. The errors of these values are less than 0.1 mmag.

A transit occurring on April 17/18 was observed from the Piskésető Station of the Konkoly Observatory. Lucky imaging technique was applied with the VersArray 1300b NTE camera on the 1m RCC telescope. In total, 4694 images were taken with 1 s exposure times. The seeing was about  $1''.5$ , and the two components could have been resolved on  $\approx 15\%$  of all images. Despite the mediocre seeing, there was chance to do photometry with image synthesis technique: the images were fitted with a superposition of two point-spread functions (PSFs). The PSF was defined for each individual images based on the observation of two bright stars in the vicinity of KOI-13. The relative position of KOI-13 A and B was kept fixed, and their joint position and the fluxes were allowed to vary. Linear resampling was applied to allow sub-pixel accuracy in the position of the model PSFs.

In the lower panel of Figure 4, we plot the result of this photometry. Small dots show the binned median of the individual flux ratios in a bin size of 27 data points. The points show the same with a bin size of 162, while error bars denote the error of the mean (i.e., standard deviation of all 162 points/ $\sqrt{162}$ ). For comparison, the predicted brightness variation is also plotted, assuming that KOI-13 A diminishes. The general similarity verifies that KOI-13 A is the one that gets fainter and thus it is the host star of KOI-13.01.

### 3.4. KOI-13.01 is a Brown Dwarf

The transit light curve has to be corrected for the light from KOI-13 B. Assuming  $\Delta V \approx 0.2$  mag, KOI-13 B contains 45% of the observed light, and the appropriate transit light curve can be derived by multiplying the excess flux by  $\approx 1.818$ . The measured depth of the transit is 4600 ppm before applying the correction, thus the corrected depth is 8400 ppm. We fitted a parametric planet model (Mandel & Agol 2002) to the reconstructed light curve via the relative size of the planet,  $r_p/r_*$ , the impact parameter,  $b$ , and the transit duration,  $D$ , assuming quadratic limb darkening with  $G_1 = 0.268$ ,  $G_2 = 0.312$  (Claret, 2003). The resulting values are  $r_p/r_* = 0.0884 \pm 0.0027$ ,  $b = 0.75$ ,  $D = 2.824 \pm 0.002$  hr. Here the errors are conservative estimates, assuming 3% ambiguity in the correction for the second light. Despite the fit is not perfect due to the prominent distortion in the light curve, the relative size of KOI-13.01 is accurate enough to determine the radius of the companion, which results to be  $2.2 \pm 0.1 R_J$ . There is no example for such large planets among the known exoplanets (their upper size limit is about  $1.7 R_J$ ; e.g., Szabó & Kiss 2011). A highly irradiated, very young ( $\approx 10$  Myr),  $0.1 M_J$  mass and coreless planet may experience this size for few million years (Fortney et al. 2007), but the derived age of KOI-13 (700 Myr) excludes such a young age. Thus, we conclude that KOI-13.01 is most likely a brown dwarf or possibly at the very low end of the red dwarf stars (Demory et al. 2009).

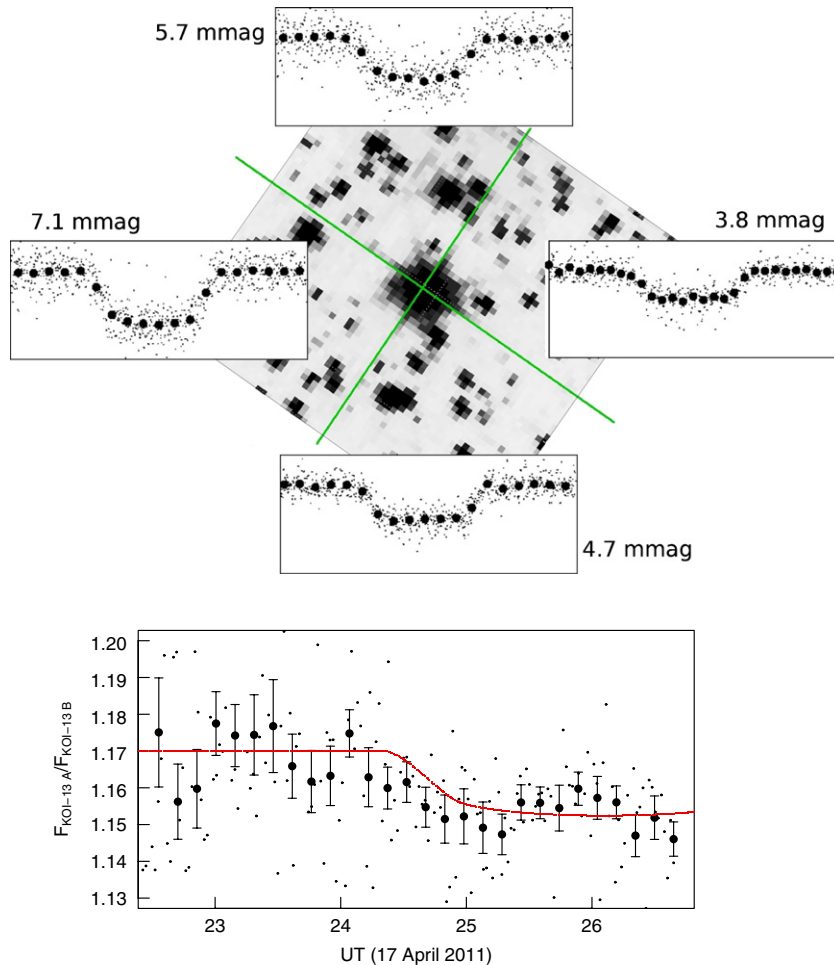
This radius is too large for a brown dwarf which is not irradiated (Baraffe et al. 2003). A red dwarf may be suggested by the observed deep secondary eclipse. The depth of the secondary eclipse is  $\approx 1/130$  of the transit, which implies a  $\approx 3150$  K effective temperature of the companion, based on our model fit in *Kepler* bandpass. This temperature is evidently too low: a red dwarf could not exhibit the observed radius unless its temperature is 3500 K ( $0.2 M_{\odot}$ , 700 Myr age).

To estimate the equilibrium temperature of the companion,  $T_{\text{eff}}^p$ , we assume uniform heat redistribution,  $T_{\text{eff}}^p = \frac{L^*}{16q^2\pi\sigma}$ . Here  $L^*$  is the luminosity of the star and  $a$  is the semi-major axis of the companion. By expressing the luminosity of the star with its  $R^*$  radius and  $T^*$  temperature, and making use of Kepler's third law, we finally get the followings:

$$\frac{T_{\text{eff}}^p}{T^*} = \sqrt{\frac{R^*}{2a}} \equiv \sqrt{\frac{R^*/R_{\odot}}{430 M^*/M_{\odot}}} \sqrt[3]{\frac{1}{P_{[\text{yr}]}}}. \quad (1)$$

The orbital period  $P$  is expressed in years. The first part takes the  $R^*/a$  term of a light curve fit as input. The right side is identical, but expressed via stellar model parameters. Substituting the parameters derived for KOI-13 A we get  $T_{\text{eff}}^p \approx 2700$  K. The estimated effective temperature exceeds the equilibrium temperature by about 17%. However, taken an uneven temperature distribution into account, a dayside temperature of  $\approx 3000$ – $3100$  K is plausible without internal heat

<sup>6</sup> MAST: <http://archive.stsci.edu/kepler/>.



**Figure 4.** Top: *Kepler* light curves from pixel-level photometry of unsaturated regions in the vicinity of KOI-13. The largest amplitude is observed to the nearest of KOI-13 A, the host star of KOI-13.01. Below: Konkoly light curve of KOI-13 A in respect to KOI-13 B on 2011 April 17. For comparison, a prediction is plotted if KOI-13 A is the host of KOI-13.01.

(A color version of this figure is available in the online journal.)

production. This makes it plausible that KOI-13.01 is a brown dwarf indeed, and its radius may be inflated simply because of the highly irradiated environment.

#### 4. DISCUSSION

KOI-13.01 is the first example of a detection of light curve distortion due to the rapid rotation of the host star. The asymmetry also implies orbital obliquity for KOI-13.01. The distortion of the transit curve is small and a unique global solution is very difficult to find, if possible at all (i.e., fitting the stellar rotation and the surface brightness distribution). Nevertheless, it is evident that the orbit is significantly oblique. Oblique orbits have been found in several exoplanetary systems, often considered as diagnostics of an orbital evolution affected by scattering processes (cf. references in Section 1). In addition to this picture, KOI-13 A has a distant stellar companion, KOI-13 B, and obviously, this system must have had a fascinating dynamical history.

KOI-13.01 is the prototype of a “hot” brown dwarf of an early-type host star, which was unprecedented before. Host stars of some exoplanets (HAT-P-7, KOI-428, WASP-33, OGLE-2) have masses between  $1.47 M_{\odot} < M_{*} < 1.52 M_{\odot}$ , predominantly F-type stars, while WASP-33 is an A5-type star (Guenther et al. 2011). In the radial velocity data, there are 24 companions be-

longing to stars with  $M_{*} > 1.8 M_{\odot}$ , and some of these objects are very probably brown dwarfs. However, these companions are not “hot,” their orbital periods span from 127 days (HD 102272) to 1191 days ( $\kappa$  CrB). It is interesting to recall the suggestion (Lovis & Mayor 2007; Bowler et al. 2010) that more massive stars tend to produce more massive planets, probably because the mass of the giant planets is scaled by the mass of the protoplanetary disk. The discovery of a close-in late-type dwarf around a massive, A-type star tends to confirm this hypothesis.

The configuration we have unraveled in KOI-13 must be rare. Searching for similar asymmetries, we tested all 474 *Kepler* candidates that had a vetting flag of 2 (i.e., passed all tests for planet nature) and which had a single exoplanet candidate companion. KOI-13 is the only clear example for such a light curve distortion. Fortunately, KOI-13 is bright enough for further observations with powerful instruments. We note that the method of symmetrized template may be too crude for the detection of fine distortions, and asymmetric analytical templates may perform better. Development of such an analytical method is in progress.

This project has been supported by the Hungarian OTKA Grants K76816, K81421, K83790 and MB08C 81013, and the “Lendület” Young Researchers’ Program of the Hungarian Academy of Sciences.

## REFERENCES

- Aitken, R. G. 1904, *Lick Obs. Bull.*, **3**, 6
- Baraffe, I., Chabrier, G., Barman, T. S., Allard, F., & Hauschildt, P. H. 2003, *A&A*, **402**, 701
- Barnes, J. W. 2009, *ApJ*, **705**, 683
- Bertelli, G., Girardi, L., Marigo, P., & Nasi, E. 2008, *A&A*, **484**, 815
- Borucki, W. J., et al. 2011, *ApJ*, submitted (arXiv:1102.0541)
- Bowler, P. B., et al. 2010, *ApJ*, **709**, 396
- Che, X., et al. 2011, *ApJ*, **732**, 68
- Ciardi, D. R., van Belle, G. T., Akeson, R. L., Thompson, R. R., Lada, E. A., & Howell, S. B. 2001, *ApJ*, **559**, 1147
- Claret, A. 2003, *A&A*, **401**, 657
- Demory, et al. 2009, *A&A*, **505**, 205
- Dommanget, J., & Nys, O. 1994, *Communications de l'Observatoire Royal de Belgique, Ser. A*, **115**, 1, (CCDM Catalog)
- Fortney, J. J., Marley, M. S., & Barnes, J. W. 2007, *ApJ*, **659**, 1661
- Gaudi, B. S., & Winn, J. N. 2007, *ApJ*, **655**, 550
- Guenther, E. W., et al. 2011, in *EPJ Web of Conferences 11, Detection and Dynamics of Transiting Exoplanets*, ed. F. Bouchy, R. Díaz, & C. Moutou (Les Ulis: EDP Sciences), 2003
- Høg, E., et al. 2000, *A&A*, **355**, L27 (Tycho-2 Catalog)
- Kharchenko, N. V., & Roeser, S. 2009, *VizieR On-line Data Catalog*, **1/280B** (Originally Published in Kharchenko, N. V. 2001, *Kinematika i Fizika Nebesnykh Tel.*, **17**, 5, 409)
- Larson, K. A., & Whittet, D. C. B. 2005, *ApJ*, **623**, 897
- Lovis, C., & Mayor, M. 2007, *A&A*, **472**, 657
- Mandel, K., & Agol, E. 2002, *ApJ*, **580**, L171
- Mason, B. D., Wycoff, G. L., Hartkopf, W. I., Douglass, G. G., & Worley, C. E. 2001, *AJ*, **122**, 3466 (WDS Catalog)
- McLaughlin, D. B. 1924, *AJ*, **36**, 16
- Monnier, J. D., et al. 2007, *Science*, **317**, 342
- Peterson, D. M., et al. 2006, *Nature*, **440**, 896
- Rossiter, R. A. 1924, *ApJ*, **60**, 15
- Shulyak, D., Tsymbal, V., Ryabchikova, T., Stütz, Ch., & Weiss, W. W. 2004, *A&A*, **428**, 993
- Szabó, Gy. M., & Kiss, L. L. 2011, *ApJ*, **727**, L44
- Szabó, Gy. M., et al. 2010, *A&A*, **523**, 84
- Tsymbal, V. 1996, in *ASP Conf. Ser. 108, M.A.S.S.; Model Atmospheres and Spectrum Synthesis*, ed. S. J. Adelman, F. Kupka, & W. W. Weiss (San Francisco, CA: ASP), 198
- von Zeipel, H. 1924, *MNRAS*, **84**, 665
- Winn, J. N., et al. 2010, *ApJ*, **718**, L145